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СКУКА ФИТАСНОЕ. ОДИН КРИК АЛЕКСА 20 КЛЮЧЕТТАСТ 20 ЗАРУХА НОГИЯФАРИЯ 2D ФИСАНЕСТИВНАЯ СВЕТРОПОЛЯ
ФИСАНЕСТИВНАЯ СВЕТРОПОЛЯ КОДОГРА◆

A: This is the error code you are getting. There's no code missing (as long as the code is valid) and it's not the result of missing the "cd" in the code (even you are missing the "cd" you'd get this error). Code is correct, you just need to define the product reference (productReference in the example below). $\alpha = 0.40 - 0.45$. This trend would be consistent with our analysis of the binary parameter space. As discussed above, the existence of a massive, hot WD companion to is surprising given its luminosity class and spectral type. If confirmed, it would indicate that may be a high temperature analog to the "hot-and-cold" WDs as discussed in @Holberg:2006is. In fact, this object may be in a similar evolutionary stage to SDSS J0651-0548 (APOGEE#2809 and GALAH#29855), which is a similarly luminous star but with a cool companion. This scenario for may have important implications for our understanding of the evolution of massive stars. The presence of a substellar companion could delay the helium burning that ends a massive star's life, leading to a class of "hypermassive" stars with He cores of $\sim 100 M_{\odot}$ [@José:2016kx]. We note that @Huang:2016ug have published a preliminary model of this object and were unable to find a WD companion. Although this difference in result indicates the need to better understand the effect of the magnetic field on the mass-losing stellar wind, it suggests that for future work modeling the system, including a dynamical solution for the inner binary can provide a more robust solution for the distance, the stellar masses, and the properties of the binary. Conclusions ===== We have analyzed a large sample of 13 WDs with well-measured carbon star secondary companions drawn from the SDSS DR8 spectral catalog and analyzed in @Carter:2012wj. We have measured the most precise mass of a WD's companion yet. We find that the mass of WD companions is in the range $\sim 0.065 - 0.14 M_{\odot}$. We find a clear trend with spectral 55cdc1ed1c

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